

Voyager Plasma and Magnetic Field Data

John Richardson

M.I.T.

Thanks to Len Burlaga for MAG data

Outline

- Overview
- Heliosheath variability
- CR-B relation
- V2 HSH evolution
- V1/V2 speed difference

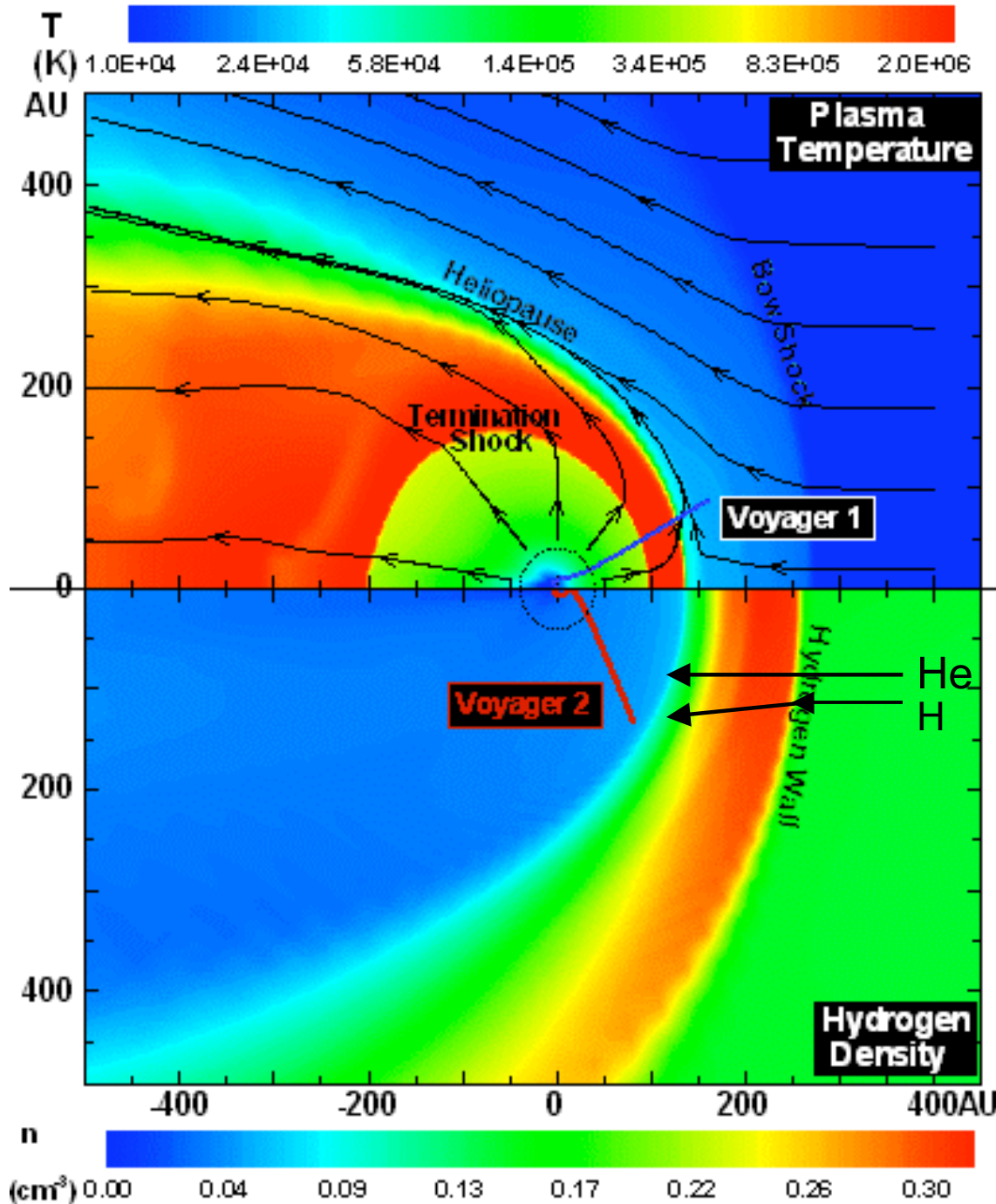
Heliosphere Overview

Asymmetric.

Thermal pressure dominated by pickup ions.

Bow shock? If LIC is supersonic, it changes to subsonic.

Mueller et al.



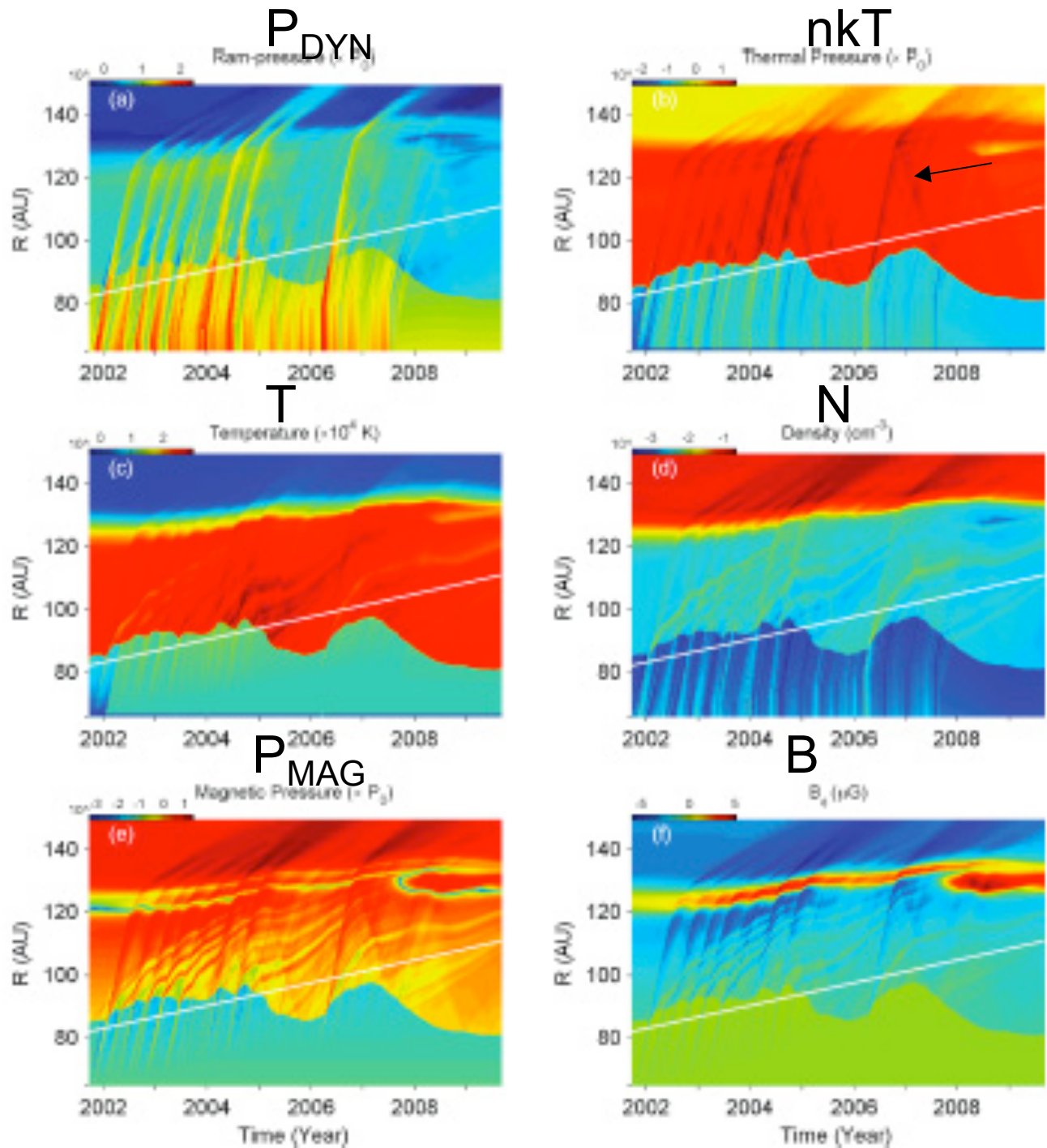
Model Prediction:
(Washimi et al.)

3-D MHD with
observed daily
average plasma
data as input.

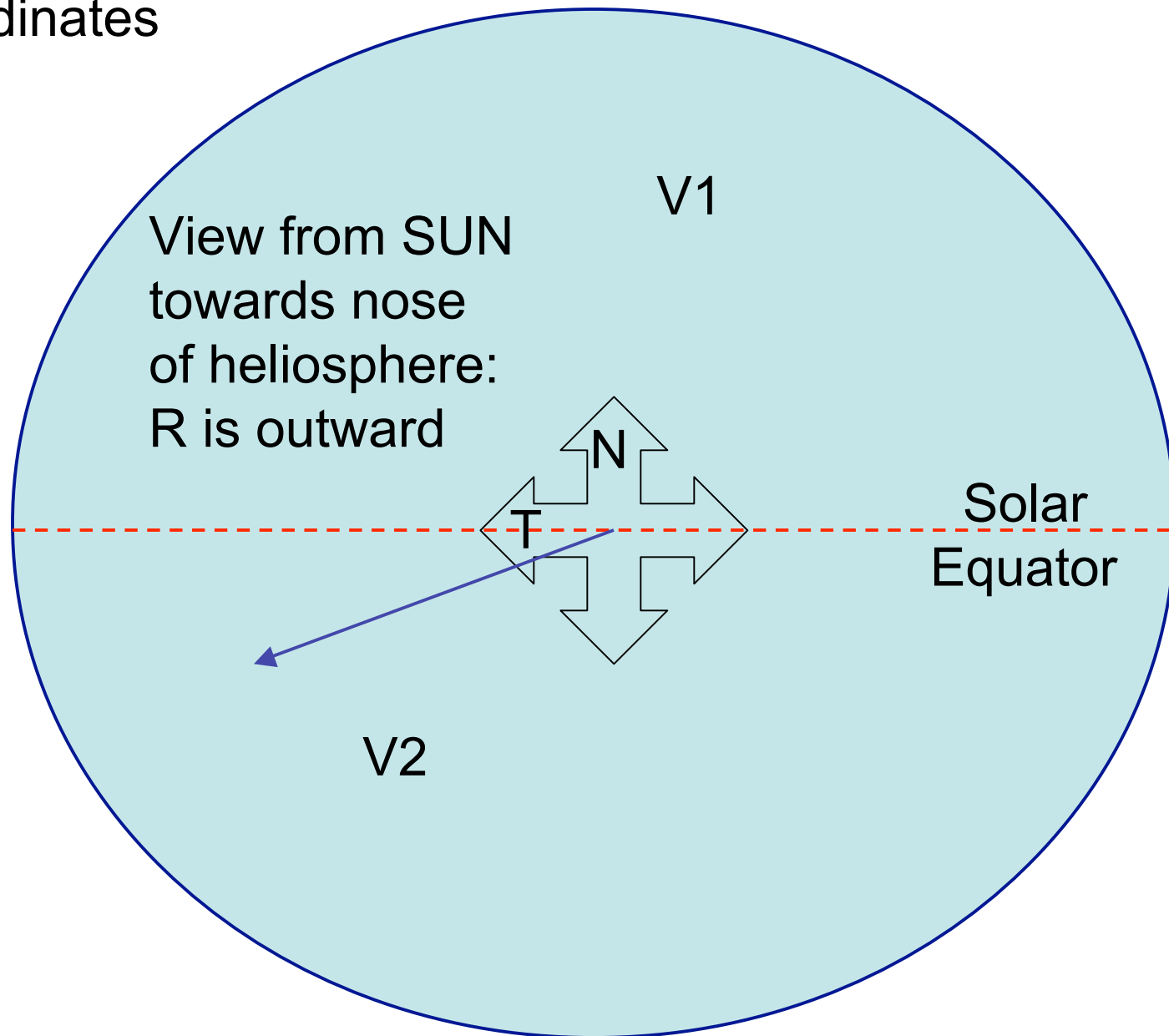
Quasi-periodic
SW pressure
pulses observed,
then propagate
through HSH.

Waves bounce
between TS and
HP.

TS motions
produce vortices



RTN coordinates



Voyager Plasma Experiment

3 Faraday Cups look sunward.

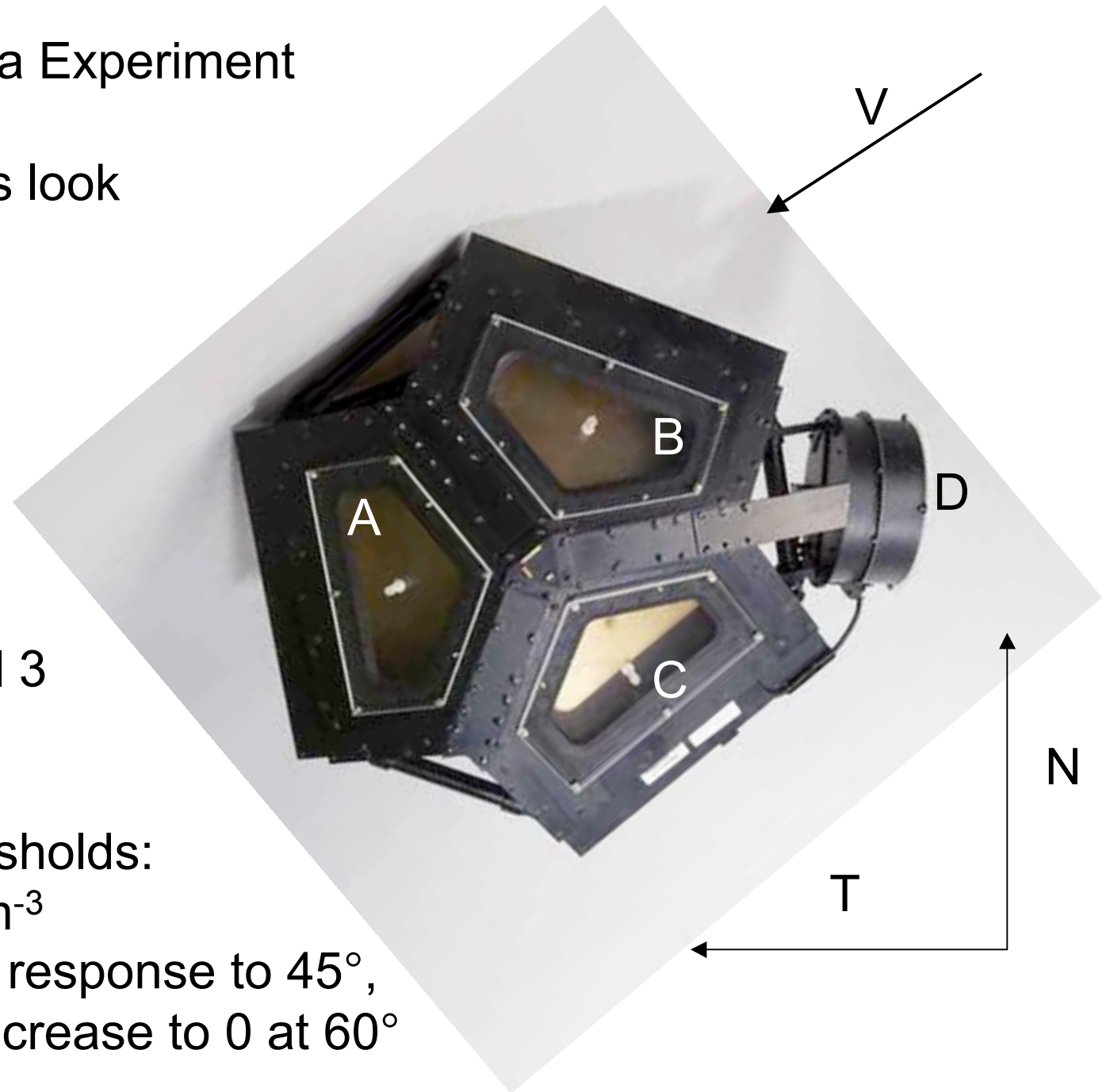
Flow is now in +R, +T, and -N direction.

Need data in all 3 Cups to fit V

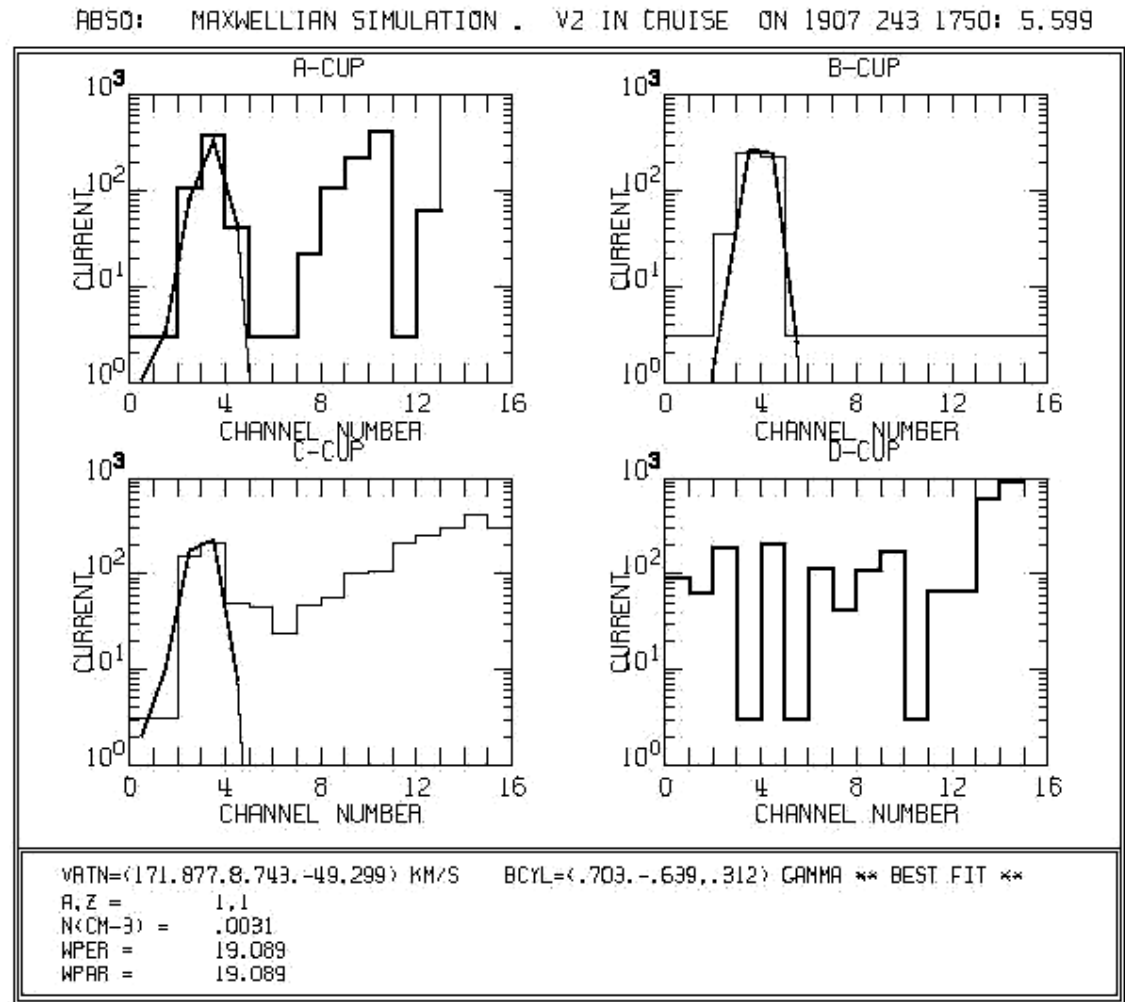
Instrument thresholds:

$N > \sim 0.0003 \text{ cm}^{-3}$

Flow angle: flat response to 45° , then linear decrease to 0 at 60°

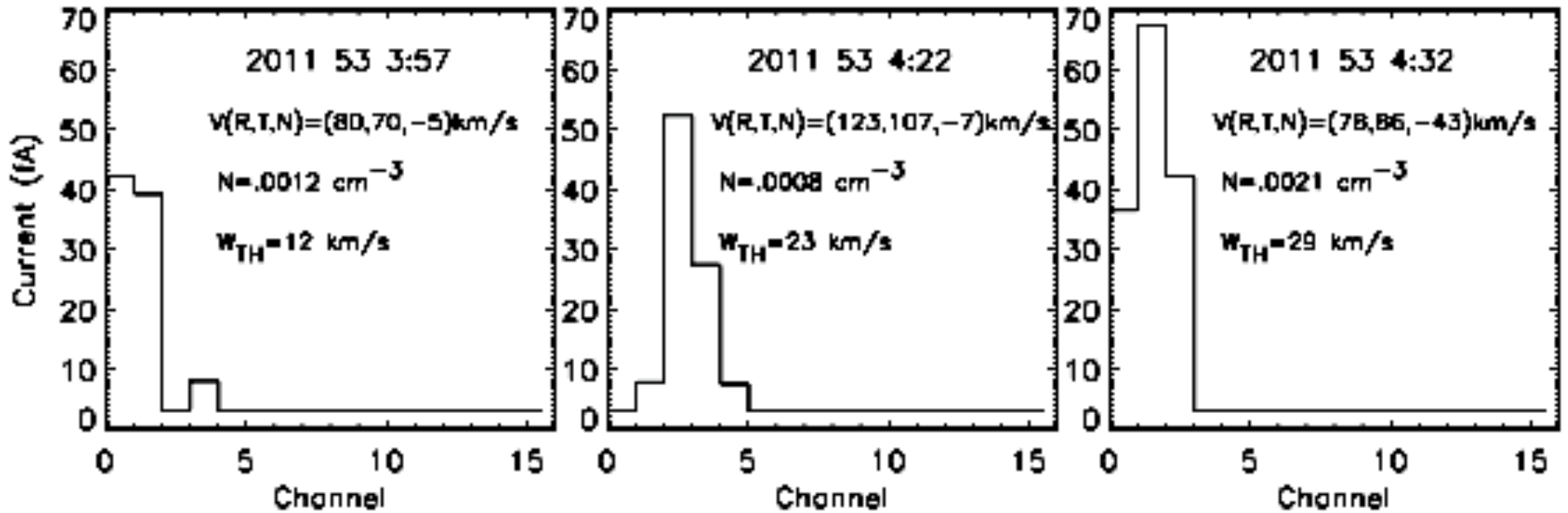


- HSH spectra
- Data in 3 cups
- Noise at higher energies
- Spectra every 192 sec.
- $V(R,T,N) =$
172, 9, -49 km/s
- $N = 0.003$ /cc
- $W_{TH} = 19.1$ km/s

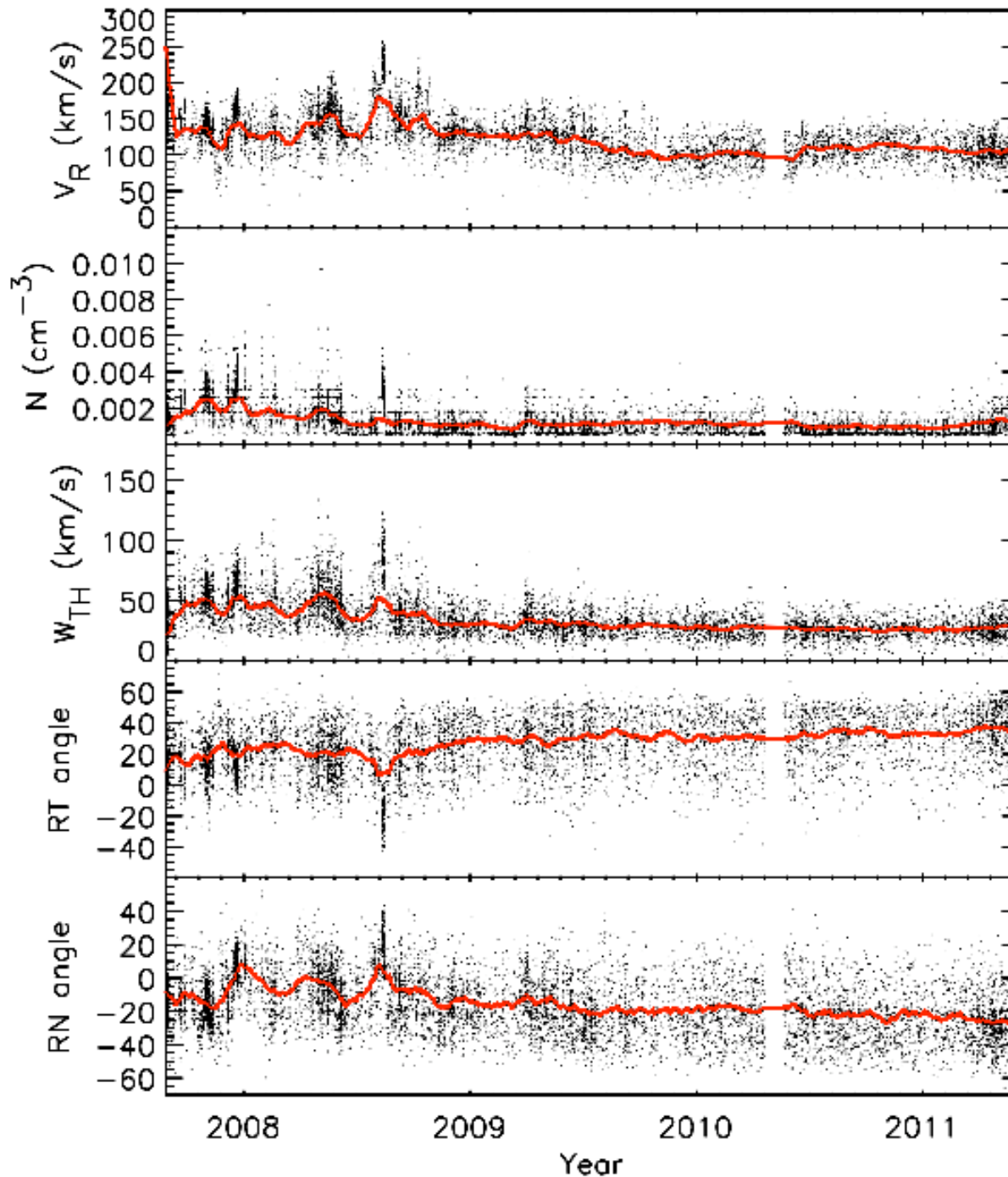


Current vs log E (10 - 5950 eV)

Plasma variability: Three spectra from within 35 minutes from B detector showing significant changes.



Current vs. log energy



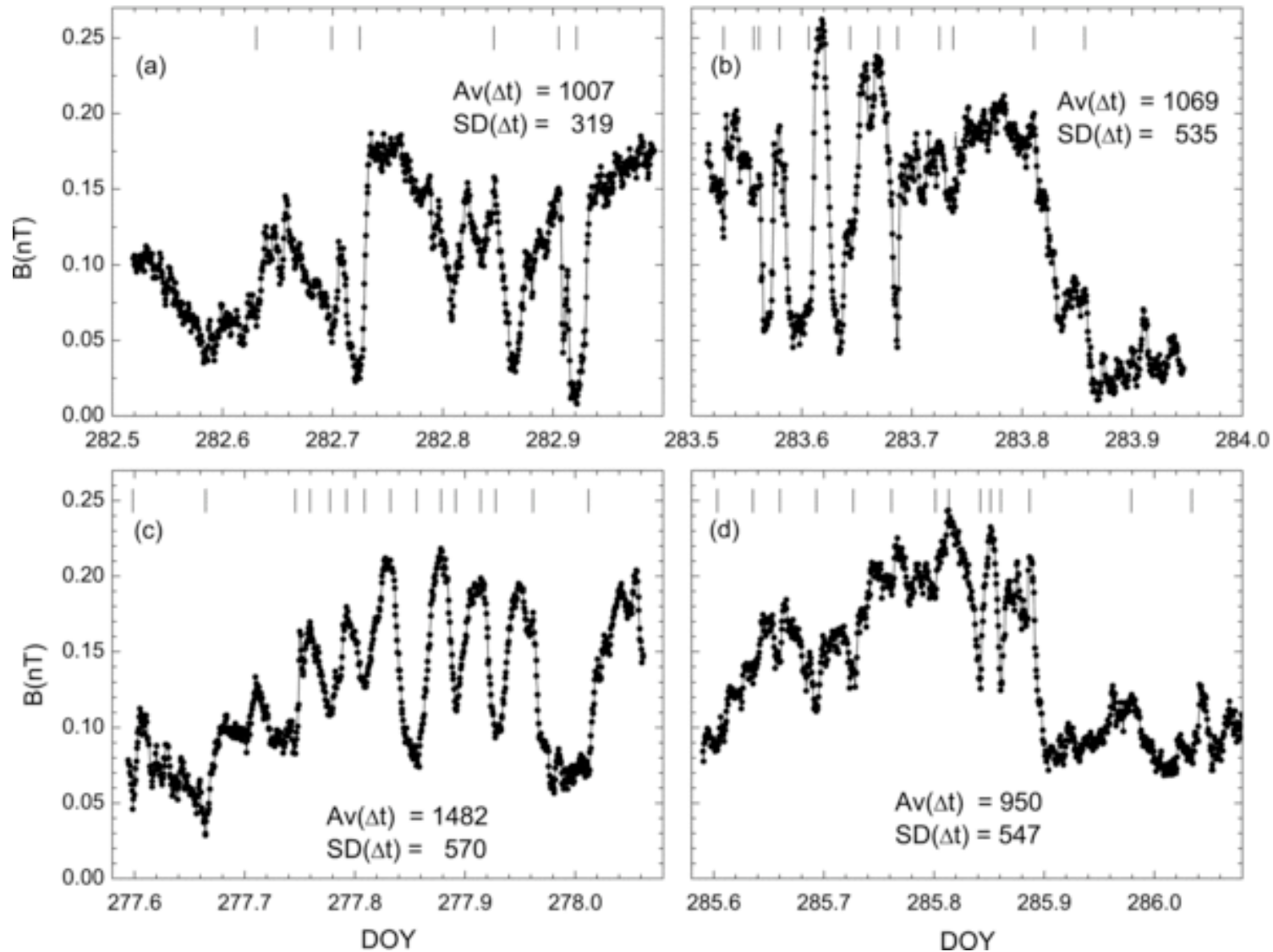
Lots of variation on short (tens of minute) time scales.

Can quantify the variability by looking at the distributions.

Distributions are generally fit well by Gaussians.

RT angle chopped off due to instrument response

Voyager 2, 2007



B variation on short time scales (48 s data, Burlaga et al., 2009)

1 keV: pui $r_l = 30,000$ km

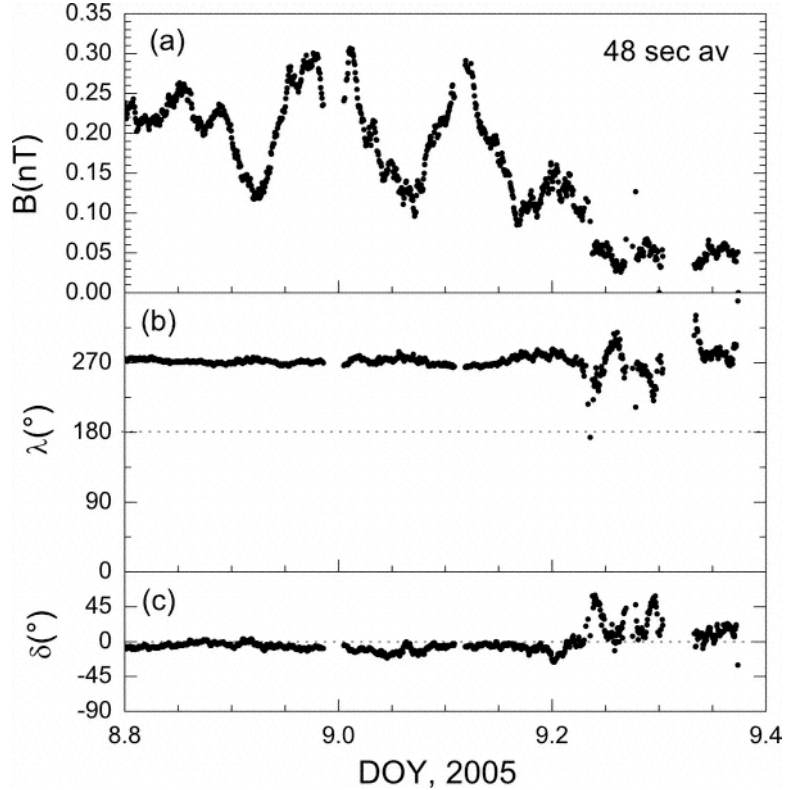
5 minutes to pass V2

4 keV: $60,000$ km

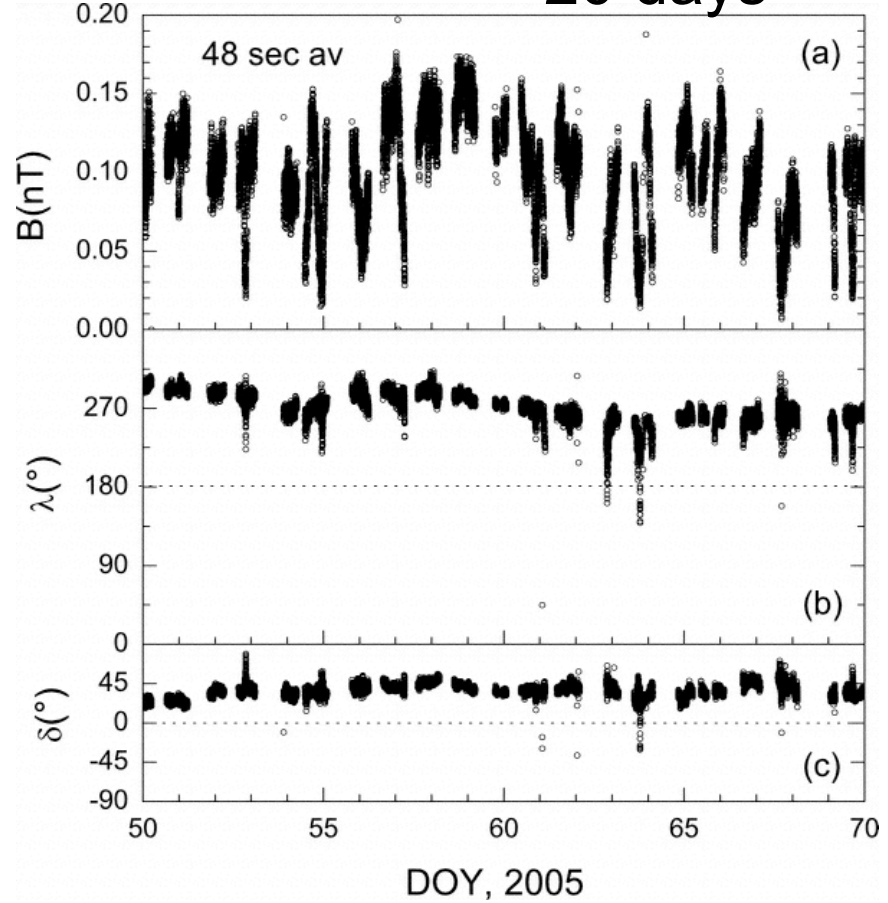
10 minutes

6 hours

Magnetic field



20 days



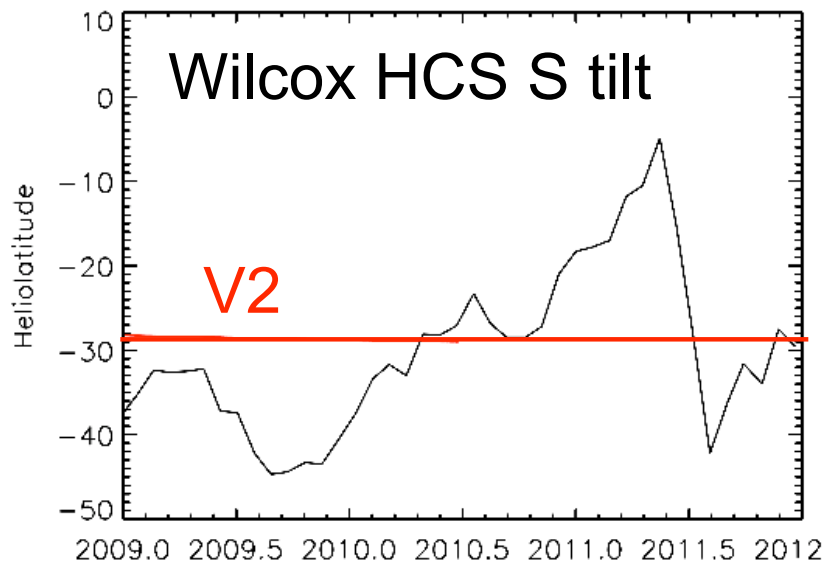
Heliosheath variability

$|B|$ highly variable

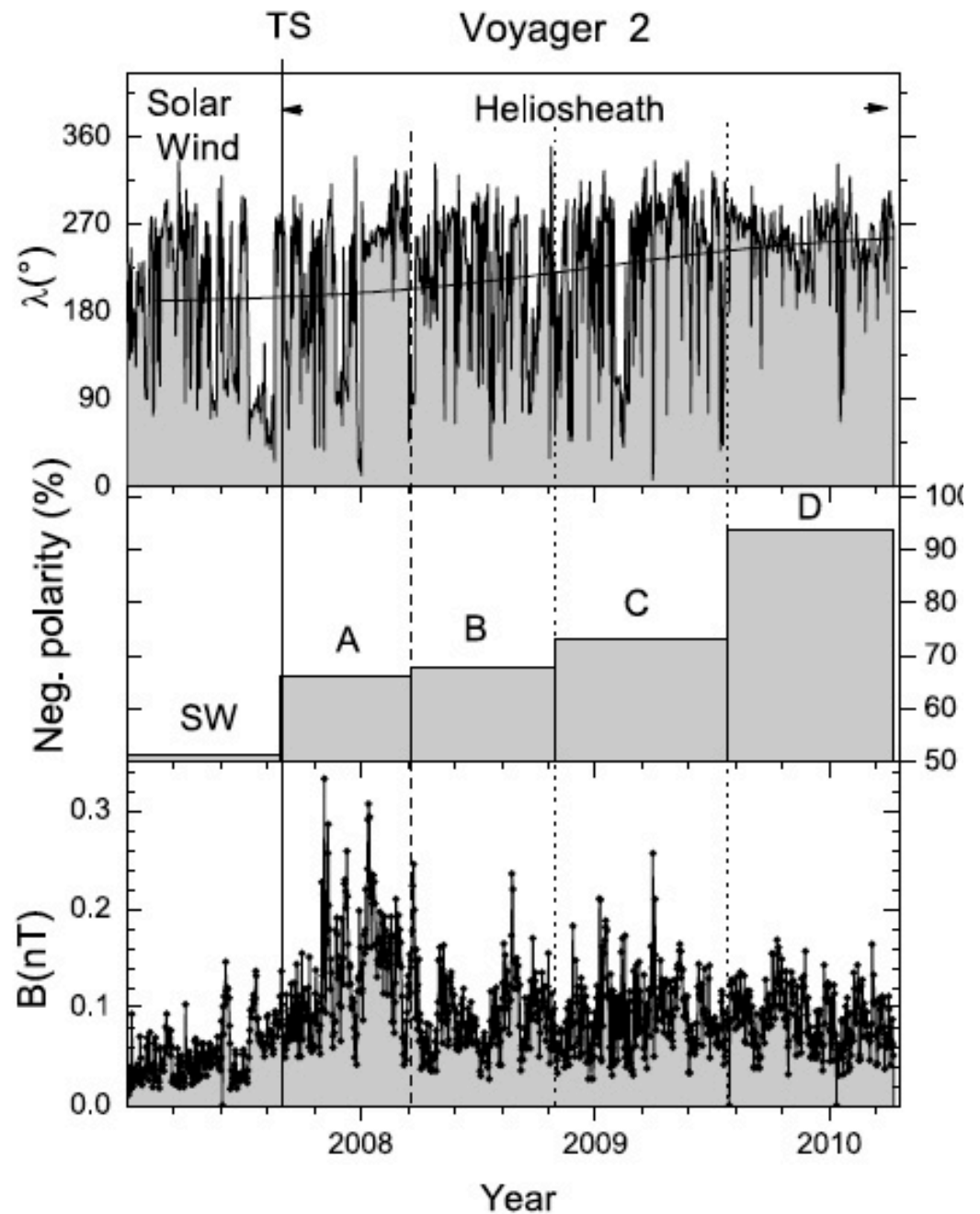
Direction \sim constant

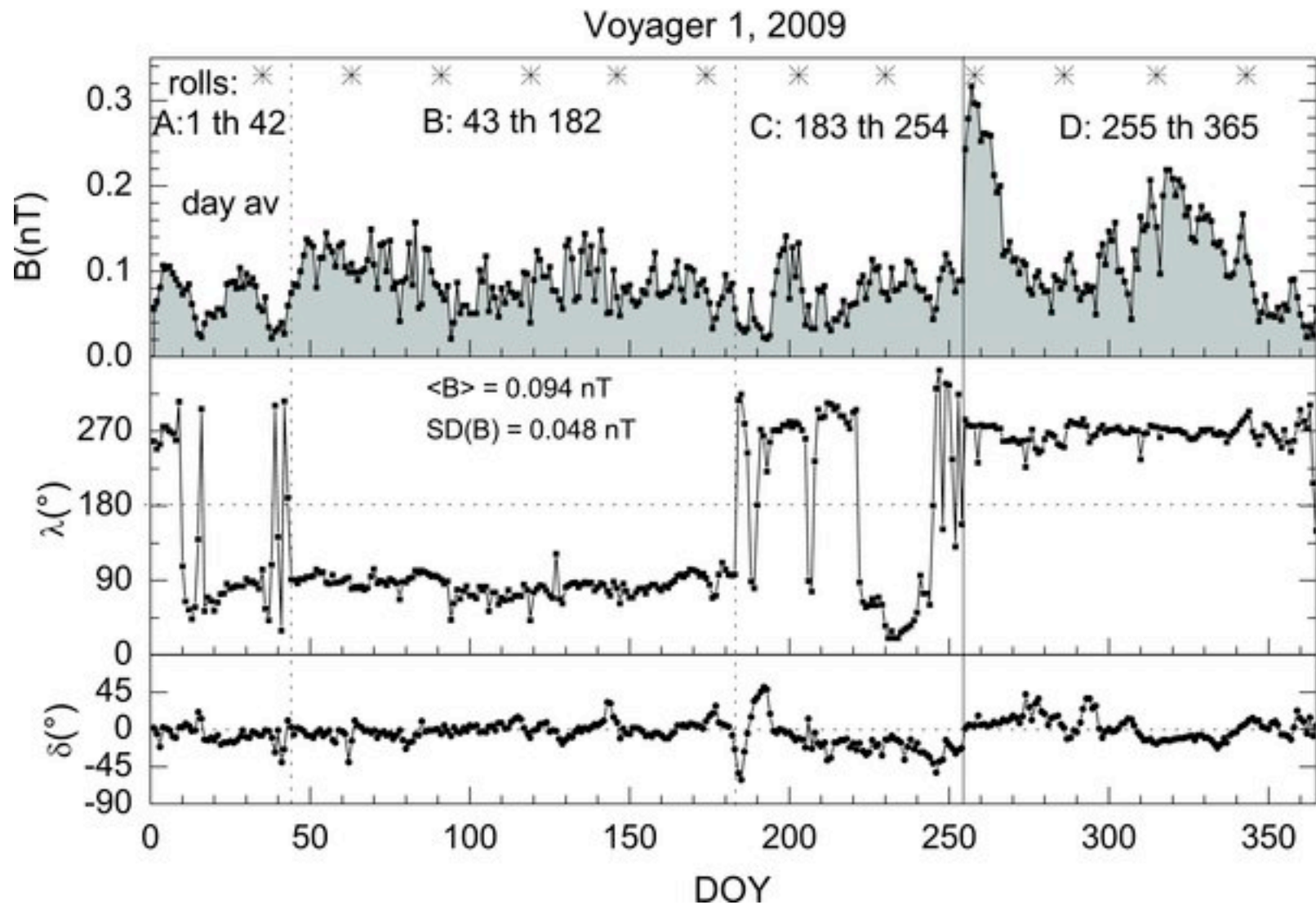
Flow is compressive

Burlaga et al., 2008



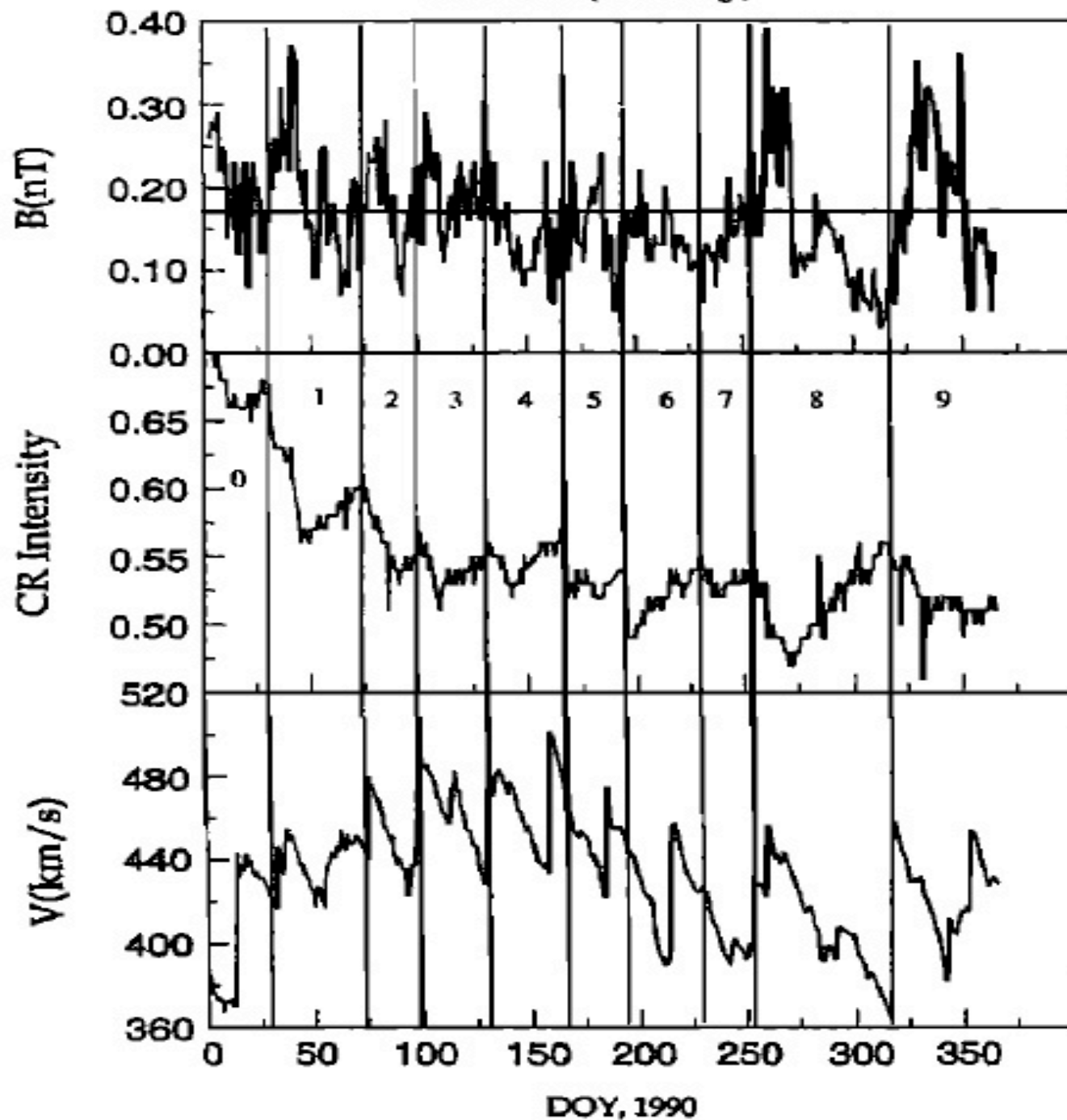
V2 almost above HCS in 2010. HCS tilt is now increasing at Sun.
 Factor of 2 changes in B from day to day common.





V1 still crossed HCS in 2009:
 low speeds = long time in sector

VOYAGER 2 (24-Hr. Avg)



CR-B relation
(Burlaga et al.)

When B is high
CRs decrease.

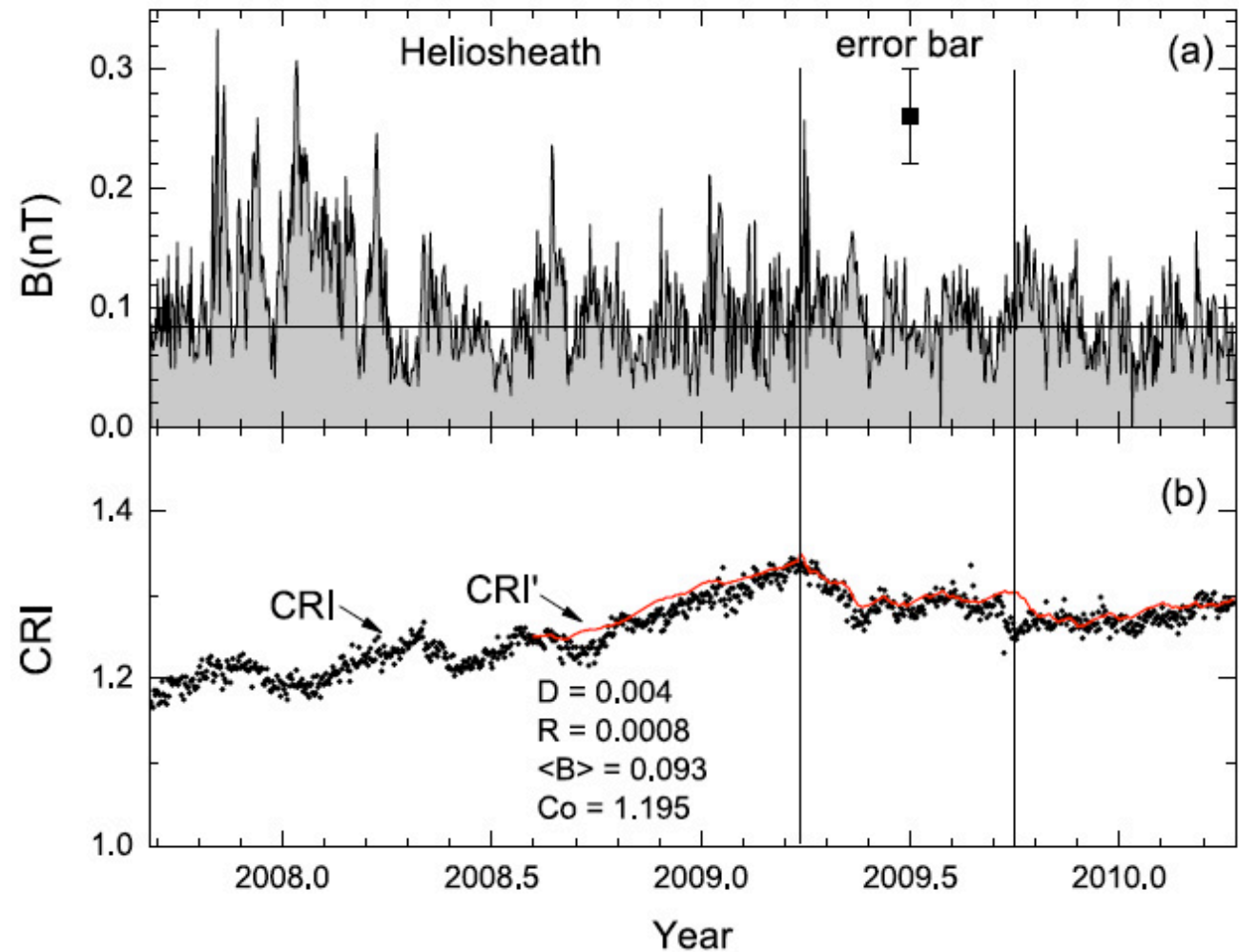
When B is low
CRs increase.

Voyager 2

CR-B fails near
TS.

Starts to work
again about
one year after
TS.

(Both at V1
and V2)

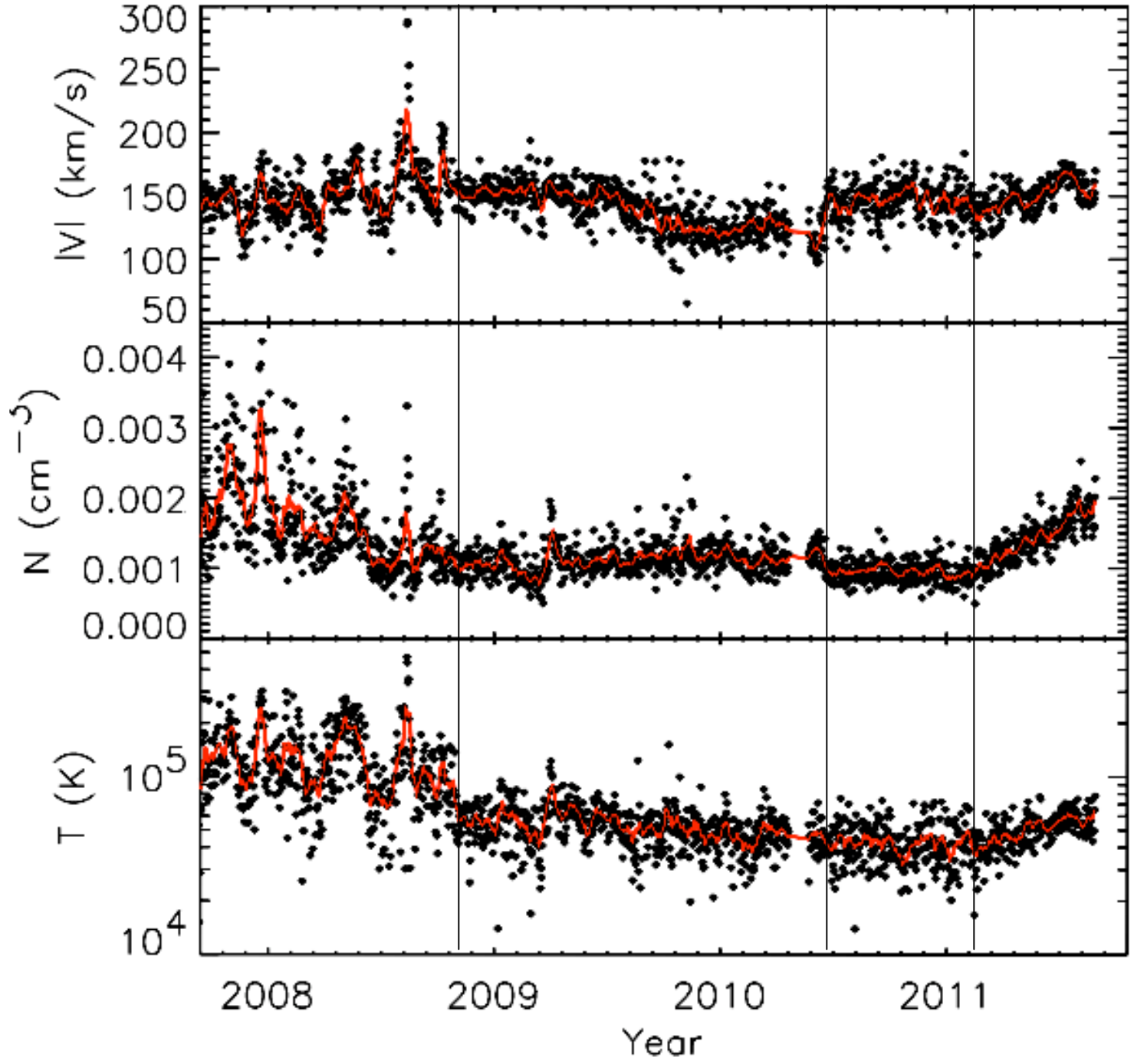


Plasma data
(1 and 11-day
averages)

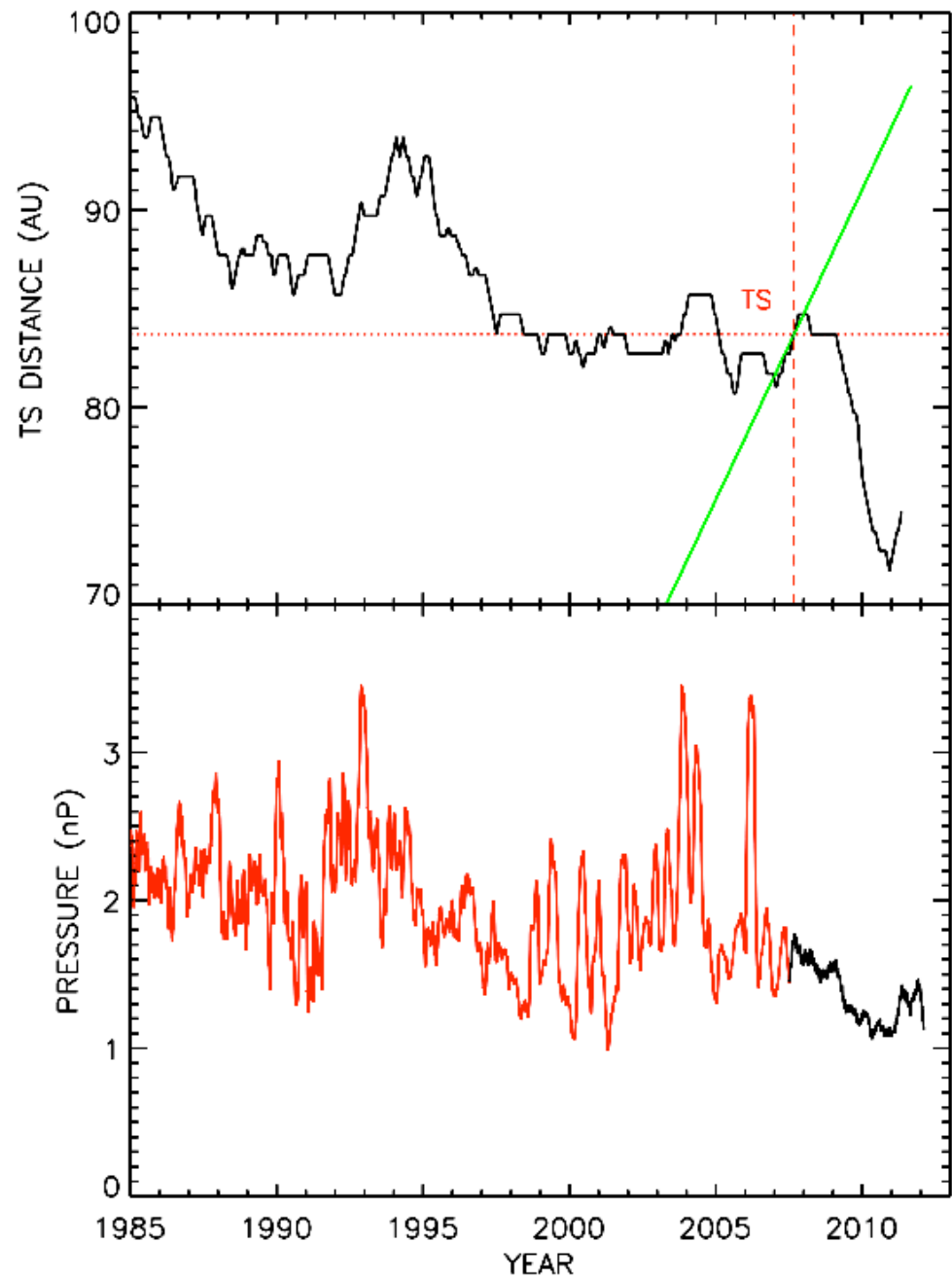
$|V| \sim \text{constant}$

N decrease
after TS, now
increasing

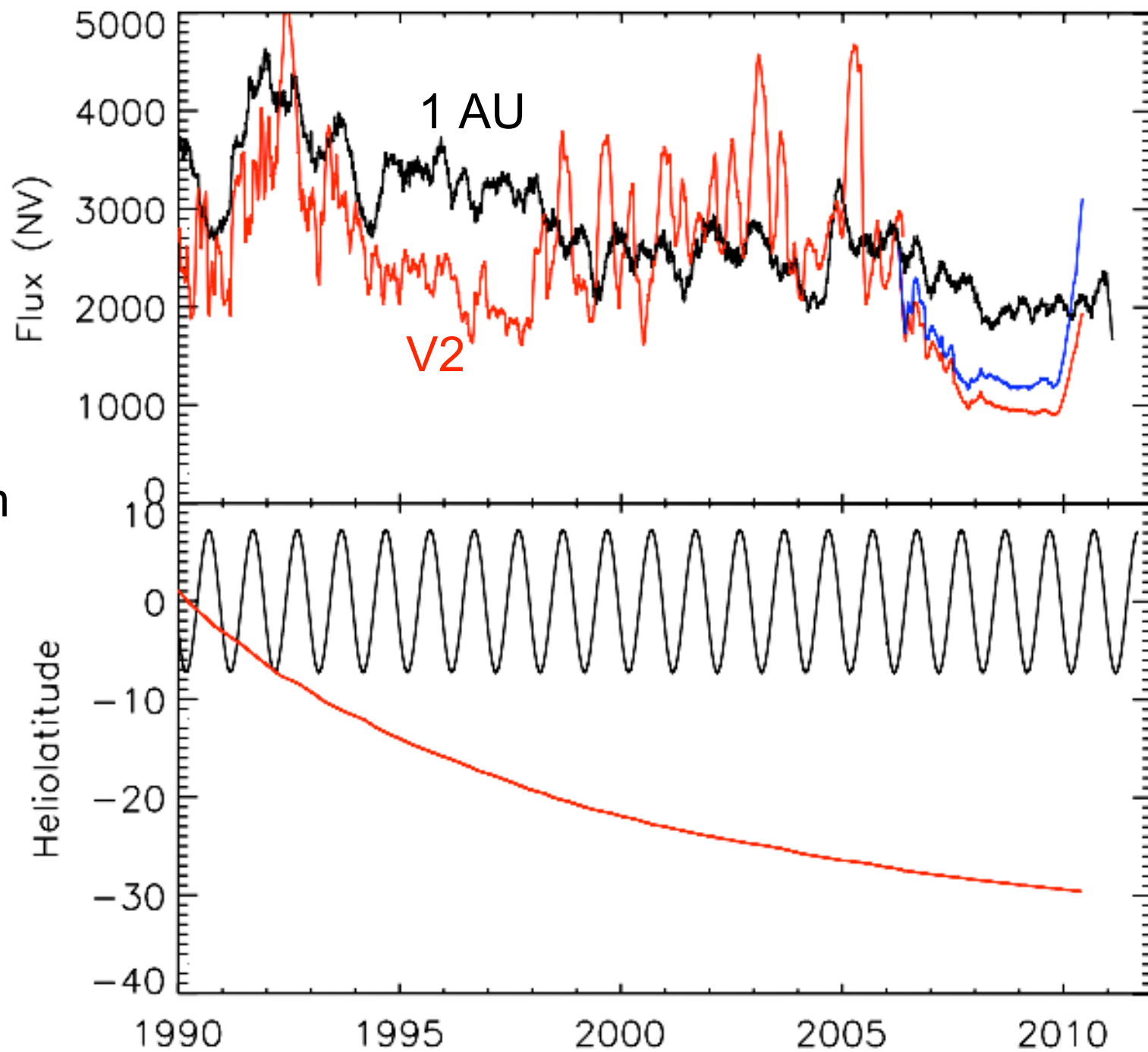
T decreased
by factor of
three, now
increasing



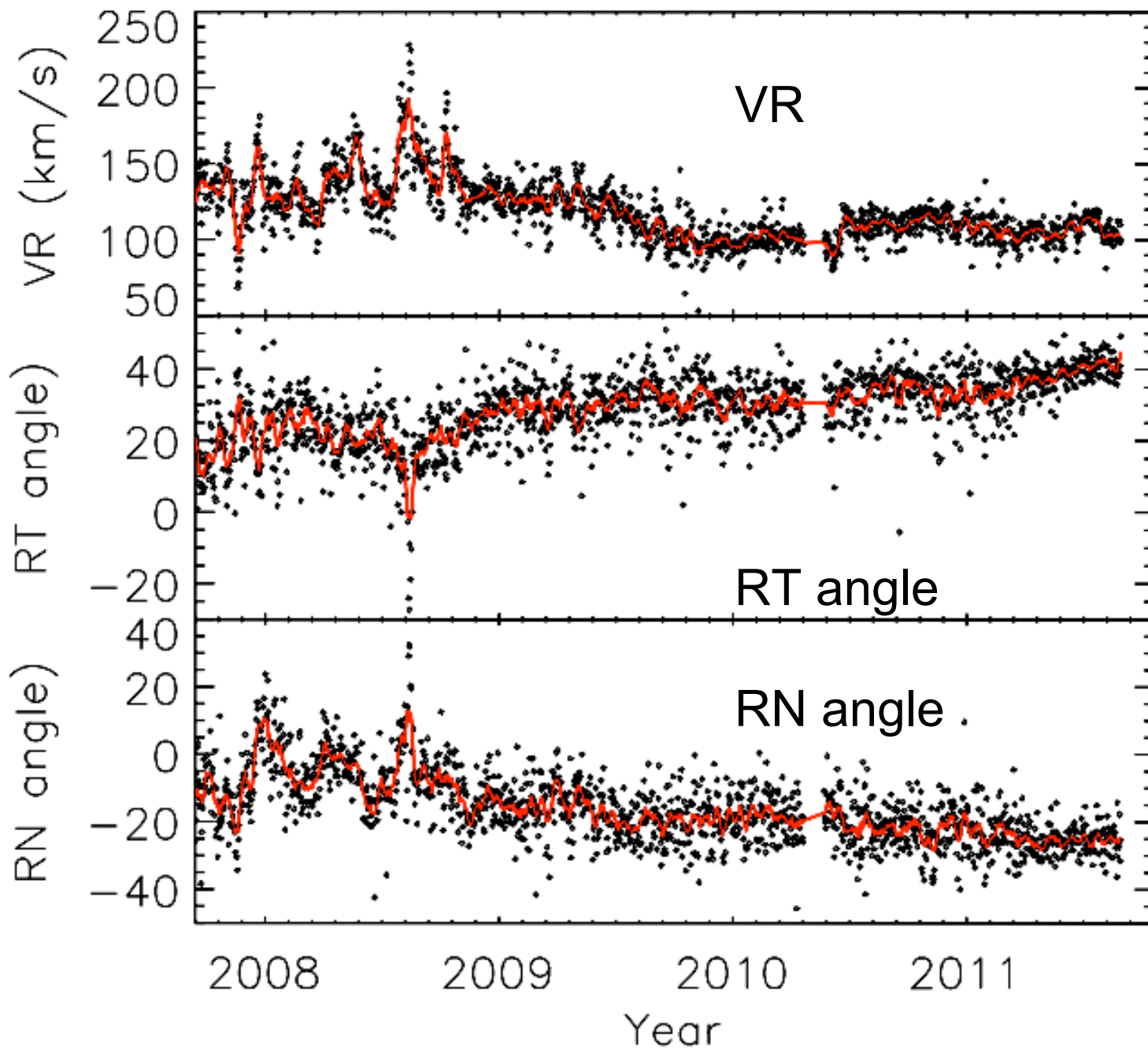
2-D HD model
with pickup
ions (Wang et al.).
V2 and 1 AU data used
as input

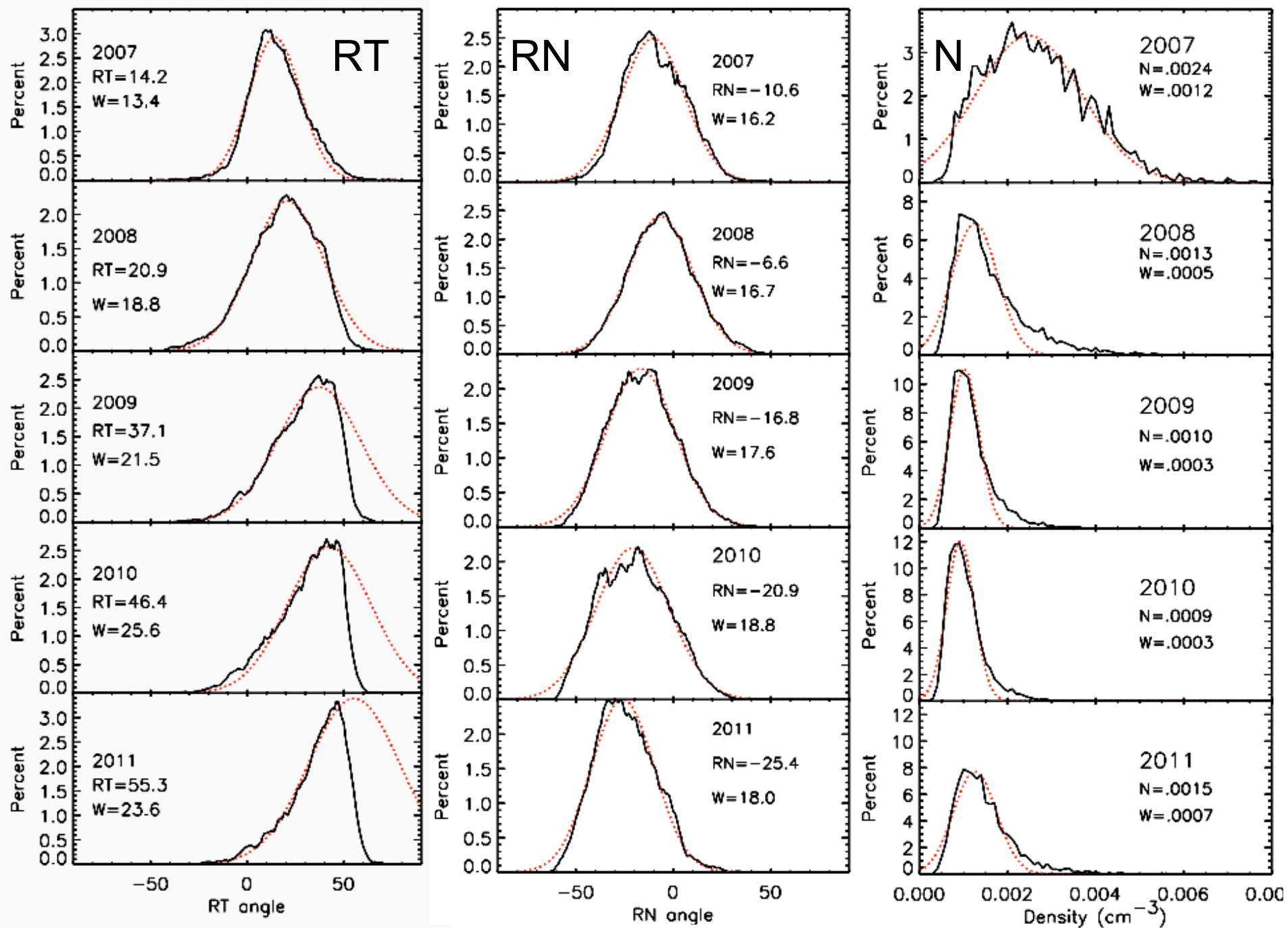


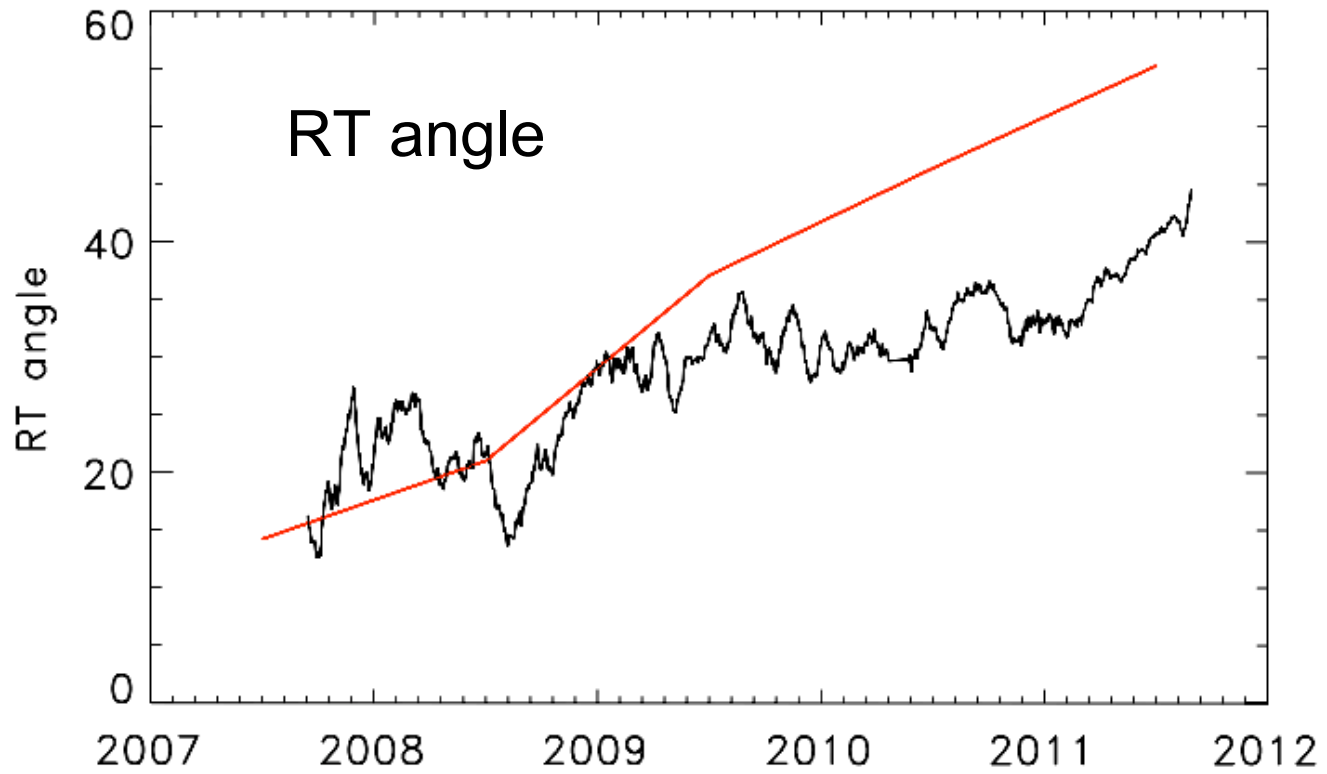
Lower
flux at
higher
latitudes
at solar
minimum



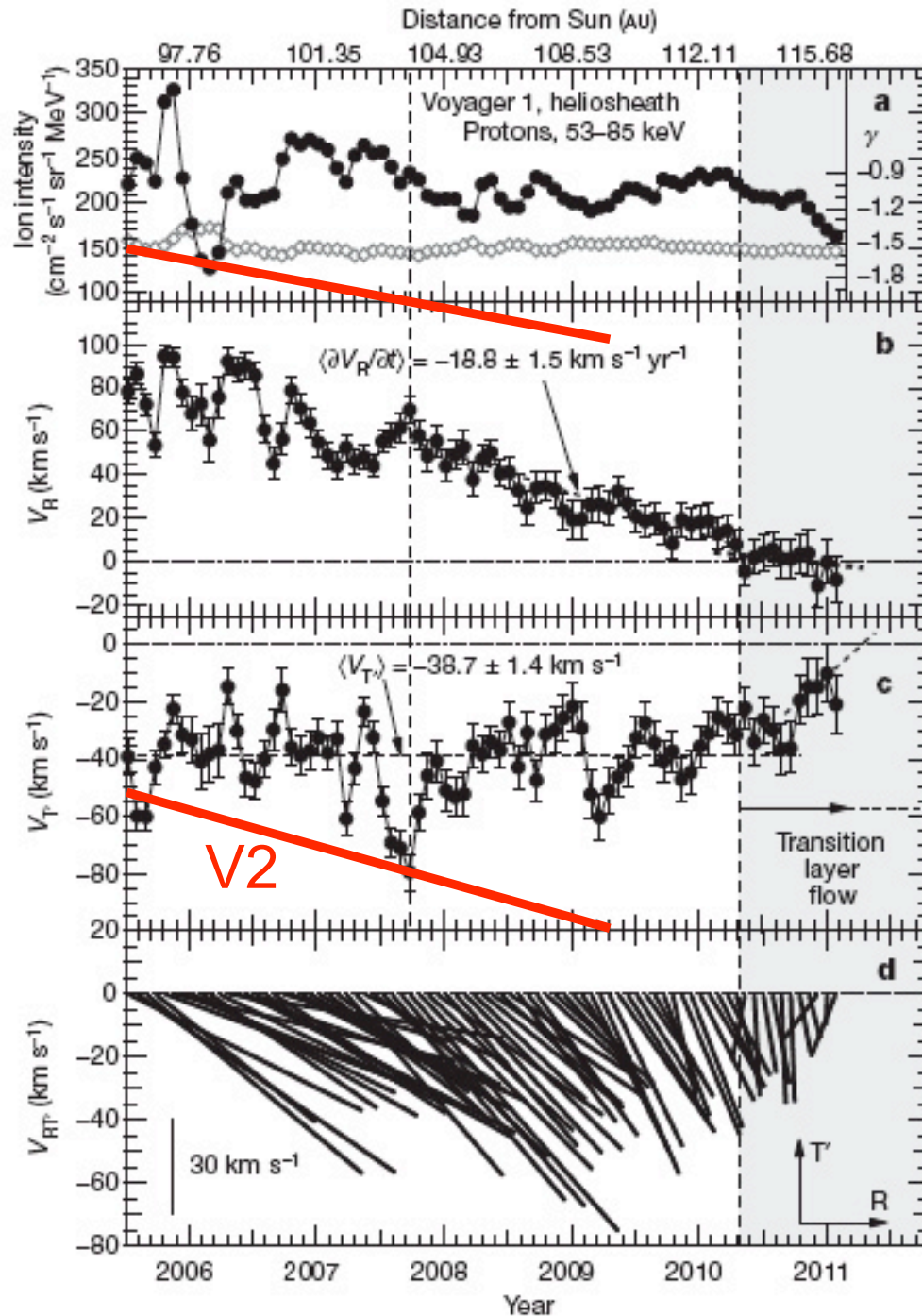
V2 1-day
And 11-day
averages







10 deg/year



Krimigis et al., Nature, 2011

V1 now at 117 AU
(V2 at 96 AU)

V1: No plasma data.

Speeds determined from
Compton-Getting effect.

V_R has gone to 0.

V_T is decreasing.

LECP does not measure V_N

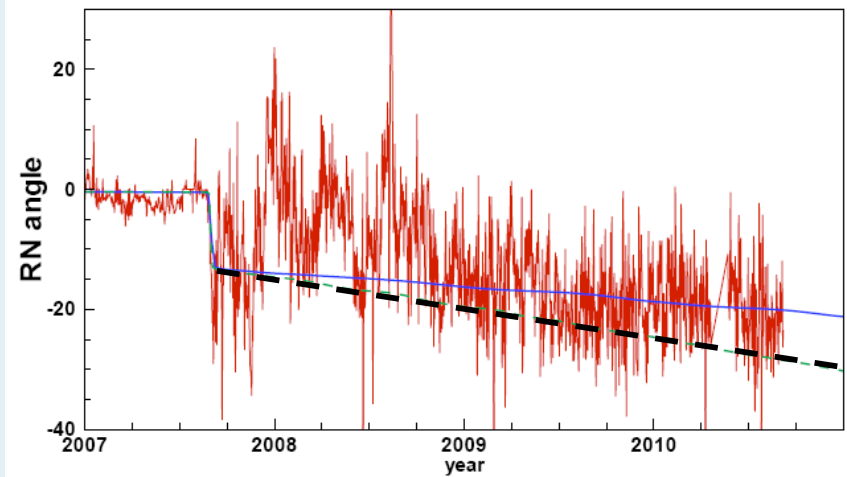
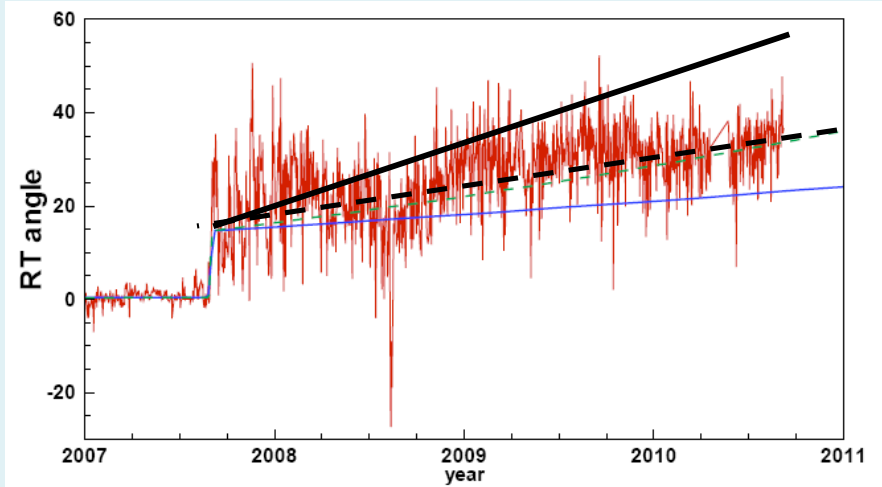
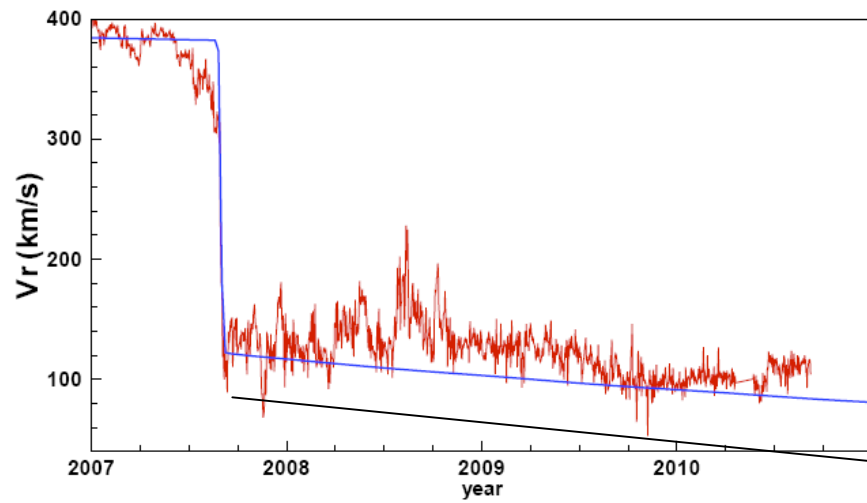
Stone et al. (2011) show V_N
also decreasing.

Ion intensity decreasing.

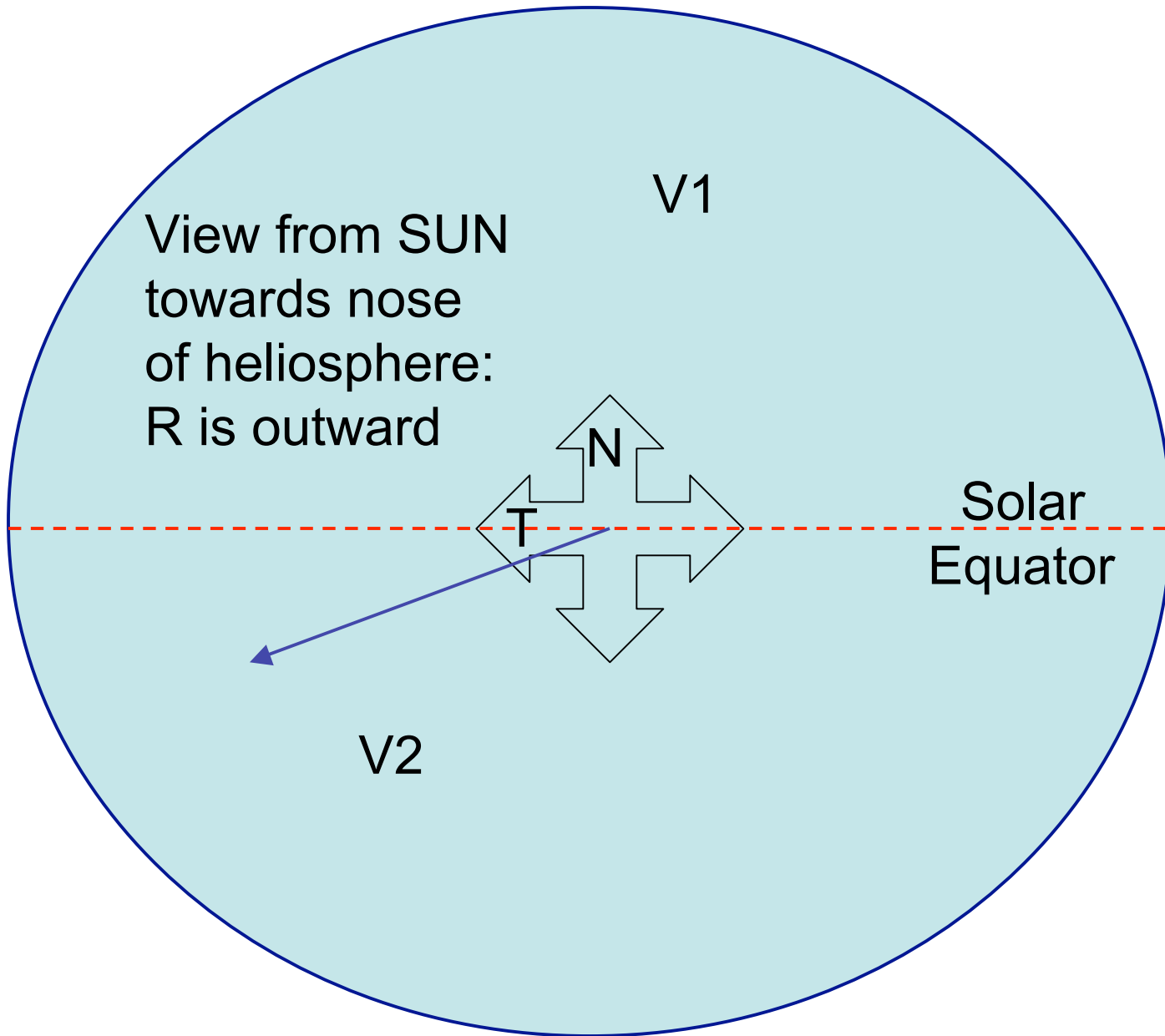
B not changing yet.

Heliopause not a sharp
boundary, but characterized
by a transition layer?

Comparison of numerical results with the Voyager 2 observations



From Borovikov et al. (2011)



Summary

- 1) HSH highly variable on time scales of tens of minutes.
- 2) Long (days) averages show coherent changes.
- 3) Recent increase in N and T
- 4) More flow deflection than model prediction

Thermal H+ and
>0.5 MeV/nuc
profiles similar.

